

Hybrid Curvatons from Broken Symmetry

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Abstract

We present a new general mechanism for generating curvature perturbations after inflation. Our model is based on the simple assumption that a field that starts to oscillate after inflation has a potential characterized by an underlying global symmetry that is slightly or badly broken. Inhomogeneous preheating occurs due to the oscillation with the broken symmetry. Unlike the traditional curvaton model, we will not identify the curvaton with the oscillating field. The curvaton is identified with the preheat field that could be either a scalar, vector, or fermionic field. We introduce an explicit mass term for the curvaton, which is important for later evolution and the decay. Our present model represents the simplest example of the hybrid of the curvatons and inhomogeneous preheating.

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1 Introduction

The traditional inflationary scenario is based on the assumption that the primordial curvature perturbation is generated by the fluctuation of the inflaton field, which leads to constraints on the traditional inflaton potential, as it is responsible for both the sufficient amount of expansion of the Universe and the generation of the curvature perturbations. Although this traditional scenario looks very simple and perhaps quite convincing, the traditional scenario sometimes requires serious fine-tunings, especially in low-scale inflationary models[1]. For example, the spectrum of the primordial curvature perturbation for the traditional scenario is given by [2]

$$\mathcal{P}_{\mathcal{R}}(k) = \frac{1}{24\pi^2 M_p^4} \frac{V_I}{\epsilon_I}, \quad (1.1)$$

where ϵ_I denotes the conventional slow-roll parameter. Obviously, fine tuning for the slow-roll parameter ϵ_I will be required for small inflationary scale. Since low-scale inflation will be very important if some gravitational effect is observed in LHC, alternatives to the traditional scenario may be desired in the near future² Recently, many alternatives to the traditional inflation have been discussed by many authors[6, 8, 9, 10, 11, 12, 13, 14, 15].³ In these alternatives the generation of the curvature perturbation is induced by the late-time conversion mechanism that characterizes the scenario. In these models both the “seed” isocurvature fluctuation is generated along with exit horizon during primordial inflation. Then the seed fluctuation is converted into curvature perturbation by the late-time mechanism. In these alternatives, and since the generation and the horizon exit of the seed fluctuations occurs during the primordial inflation, the typical length-scale is very large even though the conversion occurs very late in the scenario.

Based on recent progress in these “alternatives”, in the present paper we will consider the case where the oscillation that is associated with preheating has a potential characterized by broken global symmetry. In Sec.2 we will discuss inhomogeneous preheating

²Of course there are many other problems in low-scale inflation that are related to other cosmological issues. We think the most significant condition appears from baryogenesis[3, 4, 5] in case the electro-weak baryogenesis fails.

³The original curvaton is not suitable for low-scale inflation because of the strict lower bound for the Hubble parameter[7]. A solution to this problem is discussed by us in the first reference in Ref. [8]. Hybrid curvaton is an another solution to this problem.

with slightly broken symmetry.⁴ The oscillating field may or may not be related to the chaotic-type inflaton field. Inhomogeneous preheating with the same potential has been studied in Ref. [13] from a different viewpoint. Namely, the authors of Ref. [13] assumed that the preheat field decays instantly into light fields.⁵ The potential with broken global symmetry would be very natural, and also preheating would be very common if a field with an interaction starts to oscillate after inflation. Since our aim in this paper is to propose a simple and natural realization of the hybrid of the curvatons and inhomogeneous preheating, we will consider hybrid curvatons that work with the oscillating field accompanied by preheating, and whose potential is characterized by broken global symmetry. We will also consider late-time oscillation that induces production of the hybrid curvatons. Some variants will be discussed in the last section.

In the original scenario of the curvatons, an oscillating field (this field is identified with the curvaton in the original curvaton model) starts to dominate the energy density of the Universe because the ratio of the oscillating field to the total energy density increases during the radiation-dominated era. Then the curvatons decay to generate the cosmological perturbations. In the present paper we will consider a hybrid version of the original curvaton model, in which (unlike the traditional curvaton model) the curvaton is “not” the oscillating field. Instead of considering curvaton oscillation, we will consider generation of massive curvatons through preheating. In order to generate the initial perturbation of the curvaton number density, we will consider inhomogeneous preheating induced by a potential characterized by broken symmetry. In this case the “light field” that is associated with the primordial isocurvature perturbation is identified with the equipotential surface of the potential, which is of course not in any sense the curvatons. One might think that this scenario is not the curvatons from the viewpoint of the original discussions. However, we have already considered the basic idea in Ref. [16] where the fluctuating nucleation rate induces the fluctuation of the number density of the “topological curvatons”. We

⁴One other important idea called “inhomogeneous reheating” has been discussed by many authors[11, 10, 12]. Note that preheating is supposed to occur prior to reheating.

⁵Although the instant decay is not considered in this paper, it will be important to explain how this mechanism works in Ref. [13] and why we consider curvatons[6, 8, 16] with the same potential. See the appendix for more details and the constraints that appear for the instant-decay scenario. Note also that the authors of Ref. [13] assumed slightly broken symmetry($x \ll 1$).

have also discussed another variation in Ref. [17]. We hope there is no further confusion in the name.

Summarizing the difference between the traditional curvaton model and the hybrid version, the crucial difference is that in the original curvaton scenario the primordial isocurvature perturbation is generated for the scalar-field curvatons, while in the hybrid version the primordial perturbation does not appear for the curvatons. In the hybrid model, the fluctuation related to other light scalar fields⁶ is converted into the fluctuation of the curvaton number density after inflation. The conversion is induced by the preheating, and the preheat field is the curvaton, which could be a scalar, vector or fermionic field.

We will define a dimensionless parameter ϵ which measures the ratio between the energy stored in the oscillating field and the total energy density of the Universe. This parameter is almost unity if the oscillating field is the chaotic-type inflaton field, while it will be very small if one considers late-time oscillation. What we will consider in this paper is the oscillation and preheating with $0 < \epsilon < 1$. $\epsilon \sim 1$ corresponds to the preheating induced by the oscillation of the chaotic-type inflaton, while $\epsilon \ll 1$ is possible if the oscillating field is not associated with inflaton.

2 Slightly broken symmetry ($x \ll 1$)

The basic idea for generating the fluctuation of the preheat-field number density is not so different from the one discussed in Ref. [13], which is given in the appendix of this paper. However, as we are considering long-lived curvatons as the preheat field, there is an obvious discrepancy between the instant decay scenario[13, 14, 15] and the hybrid curvatons. The potential we will consider for the hybrid curvatons is

$$V(\phi_1, \phi_2) = \frac{m^2}{2} \left[\phi_1^2 + \frac{\phi_2^2}{1+x} \right], \quad (2.1)$$

where the dimensionless parameter x is a measure of the symmetry breaking. Due to the oscillation of the scalar field $\phi \equiv \phi_1 + i\phi_2$, the curvaton χ is produced at the ESP through preheating if there is an interaction given by $\mathcal{L} = -\frac{1}{2}g^2|\phi|^2\chi^2$. In the instant-decay scenario[13], the explicit mass term of the preheat field was not important since the

⁶This field is identified with the θ -direction in the model with slightly broken symmetry.

preheat field decays instantly when ϕ reaches the critical value, where the effective mass of the preheat field is dominated by the huge value of ϕ . However, in the present scenario we will not assume an instant-decay scenario. Instead, we will introduce an explicit mass term $\mathcal{L} = -\frac{1}{2}(m_\chi^{bare})^2\chi^2$ for the curvaton χ , which becomes important when we discuss the late-time decay of the curvaton. Then the comoving number density of the preheat field χ produced at the first scattering is given by

$$n_\chi^{(1)} = \frac{(g|\dot{\phi}_*|)^{3/2}}{(2\pi)^3} \exp \left[-\frac{\pi g(m_\chi^2 + |\phi_*|^2)}{|\dot{\phi}_*|} \right], \quad (2.2)$$

where $*$ denotes the minimum of the potential along its trajectory. The explicit form of ϕ_* and $\dot{\phi}_*$ are given in the appendix. We will consider the primordial perturbation given by $\delta\theta_0$, which leads to the fluctuation

$$\frac{\delta|\phi_*|}{|\phi_*|} = \frac{2\cos 2\theta_0}{\sin 2\theta_0} \delta\theta_0 - (\delta\theta_0)^2 \quad (2.3)$$

$$\frac{\delta|\dot{\phi}_*|}{|\dot{\phi}_*|} = -x [A\delta\theta_0 + B(\delta\theta_0)^2], \quad (2.4)$$

where A, B are defined by

$$A \equiv \frac{\sin 2\theta_0}{2(1-x\sin^2\theta_0)} \simeq \frac{\sin 2\theta_0}{2} + \mathcal{O}(x) \quad (2.5)$$

$$B \equiv \frac{4(1-x\sin^2\theta_0)\cos 2\theta_0 + x\sin^2 2\theta_0}{8(1-x\sin^2\theta_0)^2} \simeq \frac{\cos 2\theta_0}{2} + \mathcal{O}(x). \quad (2.6)$$

Then the fluctuation of the preheat-field number density is

$$\begin{aligned} \frac{\delta n_\chi}{n_\chi} &\simeq \left(\frac{3}{2} + \frac{\pi g|\phi_*|^2}{|\dot{\phi}_*|} \right) \frac{\delta|\dot{\phi}_*|}{|\dot{\phi}_*|} - \frac{2\pi g|\phi_*|^2}{|\dot{\phi}_*|} \frac{\delta|\phi_*|}{\phi_*} \\ &\simeq -x \left(\frac{3}{2} + \mathcal{O}(x^2) \right) [A\delta\theta_0 + B(\delta\theta_0)^2] - (\mathcal{O}(x^2)) \left[\frac{2\cos 2\theta_0}{\sin 2\theta_0} \delta\theta_0 - (\delta\theta_0)^2 \right] \\ &\simeq -\frac{3x}{2} [A\delta\theta_0 + B(\delta\theta_0)^2] + \mathcal{O}(x^2). \end{aligned} \quad (2.7)$$

In the instant-decay scenario, the preheat field decays instantaneously when ϕ reaches its maximum value just after the first scattering. Then, assuming that the cosmological perturbation is generated by the instant decay, the decay products must nearly dominate the energy of the Universe. This requirement puts a crucial bound on the dimensionless parameter ϵ . Since the energy density of the preheat field is bounded above by the energy of the oscillating field, the requirement for the preheat-field domination leads to $\epsilon \simeq 1$.

The condition $\epsilon \simeq 1$ suggests that the oscillating field must be the inflaton of chaotic inflation, which finally leads to the bounds that we have obtained in the appendix. On the other hand, $\epsilon \ll 1$ is acceptable for the curvatons since the ratio of the curvaton to the total energy density will grow during the radiation-dominated era.

In the hybrid curvaton model, the condition from the COBE normalization is given by

$$\left[\frac{3x \sin 2\theta_0}{4} \right] \delta\theta_0 \sim 10^{-5}, \quad (2.8)$$

which comes from Eq.(2.7). Here we will introduce a new parameter ϕ_{osc} which denotes the initial value of the oscillating field at the time when the oscillation starts. Note that $\phi_{osc} \sim M_p$ in the instant-decay scenario while $\phi_{osc} \ll M_p$ is allowed in the curvaton scenario. The fluctuation $\delta\theta_0$ is determined by the value ϕ_I during horizon exit and the Hubble parameter H_I during inflation, and is given by $\delta\theta_0 \simeq H_I/2\pi\phi_I$. This value becomes $\delta\theta_0 \sim H_I/20\pi M_p$ in the instant-decay scenario. On the other hand, in the curvaton scenario it is possible to have either $\phi_I \gg \phi_{osc}$ or $\phi_I \ll \phi_{osc}$ since ϕ is not necessarily the inflaton. Moreover, $\delta\theta_0$ may evolve after the inflationary stage until the time when the curvaton is generated through preheating. Hence, the fluctuation $\delta\theta_0$ is not a parameter that is completely determined by the inflationary model.

As we are considering long-lived curvatons, we cannot disregard the confining potential that is induced as the backreaction of the preheat field, and also the fluctuations that may be induced at the second and the third scatterings. After the first scattering that we have discussed above, the scatterings occur successively where the number densities are given by

$$n_k^{j+1} = b_k^j n_k^j. \quad (2.9)$$

Here the coefficient b is given by

$$b_k^j = 1 + 2e^{-\pi\mu^2} - 2 \sin \theta^j e^{-\pi\mu^2/2} \sqrt{1 + e^{-\pi\mu^2}}, \quad (2.10)$$

where $\mu^2 = (k^2 + (m_\chi^{eff})^2)/g|\dot{\phi}|$ and θ^j is a relative phase. Of course we cannot disregard the fluctuations coming from δb_k^j , although it will be very difficult to say which scattering induces the most significant effect on the fluctuation. In Fig.1 we show an obvious example adding an effective mass $m_\chi^{eff} = m_\chi^{bare} + g\phi$, where it will be very difficult to examine the generic properties of the later scatterings. Instead of solving the process that is

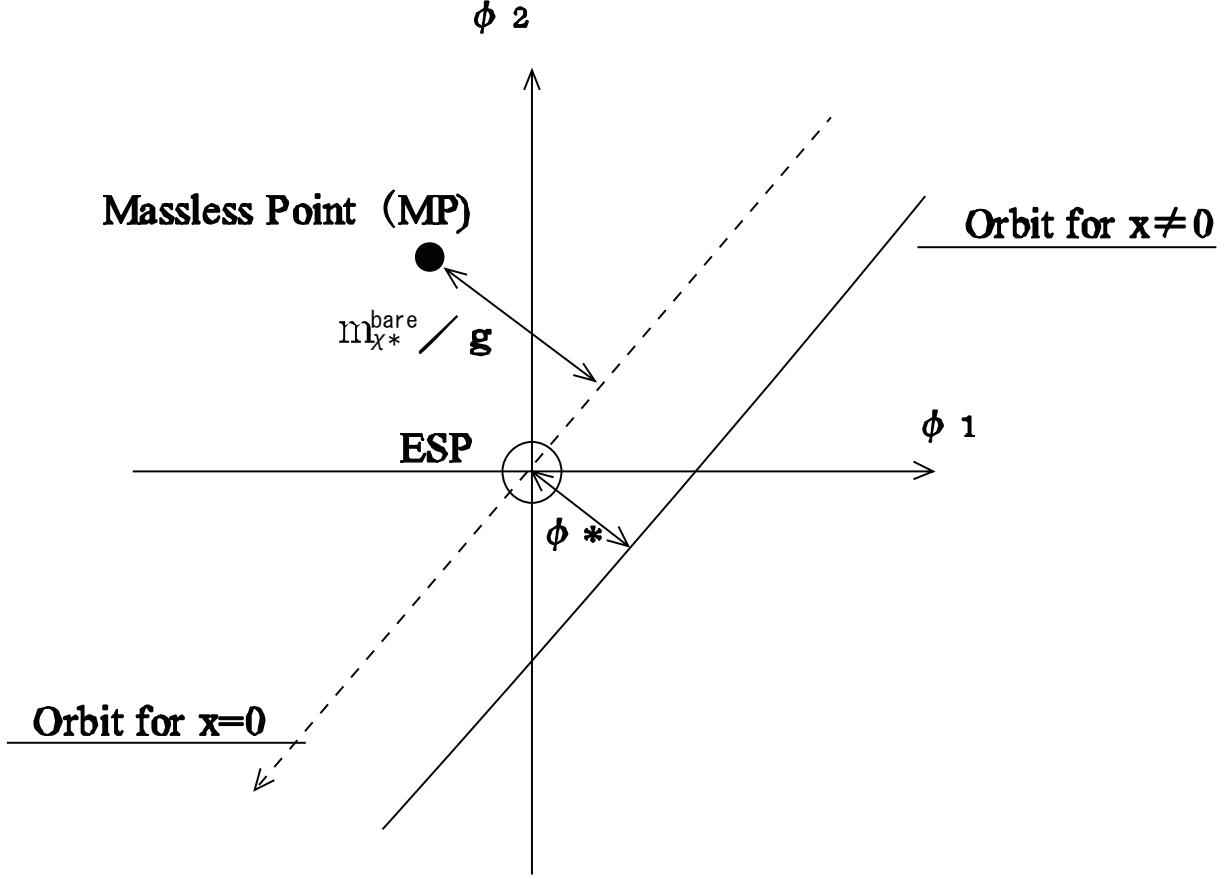


Figure 1: Including $m_\chi^{bare} \neq 0$, the massless point (MP) where preheat field becomes massless is shifted from the ESP. Because of the shift denoted by m_χ^{bare} , the minimum value of the preheat-field mass is given by $|m_\chi| = g\phi_* + m_\chi^{bare}$.

not essential for the argument, we will simply assume that the first scattering gives the dominant contribution to the fluctuation. Note that even if the later scattering happens to generate a comparable magnitude of fluctuation, the basic form of the result will not be altered very much.

Initially the curvaton energy density is $\rho_c = m_\chi n_\chi^{ini} \simeq (g\phi)(g\dot{\phi}_*)^{3/2}/(2\pi)^3 \leq (\dot{\phi}_*)^2 \simeq m^2\phi_{osc}^2$, while the total energy density of the Universe is $\rho_{tot} \simeq H_{osc}^2 M_p^2 \simeq m^2 M_p^2$, thus we obtain the initial ratio of the energy density $r \equiv \rho_c/\rho_{tot} \leq \phi_{osc}^2/M_p^2$, which equation is similar to that found in the conventional curvatons. The late-time evolution of the dimensionless parameter r is also the same as the conventional curvatons, but ρ_χ evolves as $\rho_\chi \propto a^{-4}$ [18] until m_χ^{bare} dominates m_χ . The only and the crucial difference appears from the condition coming from the observed curvature perturbation. In the conventional

curvatons, this condition is given by

$$\mathcal{P}_\zeta \simeq \frac{rH_I}{2\pi\sigma_{osc}} \sim 10^{-5}, \quad (2.11)$$

where σ_{osc} denotes the value of the curvaton σ at the onset of the oscillation. This condition does not appear in our hybrid-type curvaton model, since in our model the fluctuation of the curvaton number density is “not” related to the primordial isocurvature perturbation that might be generated for the curvaton χ , but is related to the parameter $\delta\theta_0$ of the oscillating field ϕ . Of course the oscillating field ϕ is not the curvaton. In the conventional curvaton model, the decay rate of the curvaton is supposed to be not much smaller than $\Gamma \sim m^3/M_p$, thus the mass of the oscillating potential (m) puts a lower bound on the decay temperature T_d . The bound from the decay temperature is of course important in the present model. However, unlike the conventional curvatons, the decay rate of the hybrid-type curvaton is not determined by the mass m of the oscillating field, but is determined by m_χ and its coupling g_d to light field. Therefore, in the hybrid-type curvaton the mass of the oscillating field(m) that determines the initial value of the total energy density through $H_{osc} \simeq m$ is not relevant to the decay temperature T_d , which removes one of the most important conditions in the conventional curvatons.

In the hybrid-type scenario, difficulty may appear from the condition related to the dimensionless parameter x . From Eq.(A.9) in the appendix, the condition needed for successful preheating is

$$x^2 < \frac{m}{g\phi_{osc} \sin^2 2\theta_0}. \quad (2.12)$$

From Eq.(2.8), we can see the condition for the curvature perturbation

$$x \sim \frac{4}{3 \sin 2\theta_0} \frac{10^{-5}}{\delta\theta_0}. \quad (2.13)$$

Finally, the non-Gaussian condition $f_{NL} < 100$ is given by

$$x > \frac{1}{100} \left(\frac{\cos 2\theta_0}{\sin^2 2\theta_0} \right). \quad (2.14)$$

From Eq.(2.12) and Eq.(2.13), we obtain

$$\delta\theta_0 > 10^{-5} \sqrt{\frac{g\phi_{osc}}{m}}. \quad (2.15)$$

From Eq.(2.14) and Eq.(2.13) we obtain

$$\delta\theta_0 < \frac{\sin 2\theta_0}{\cos 2\theta_0} 10^{-3}. \quad (2.16)$$

Combining the last condition and Eq.(2.15), we find the condition for the ratio ϕ_{osc}/m , which is given by

$$\frac{\phi_{osc}}{m} < \left(\frac{\sin 2\theta_0}{\cos 2\theta_0} \right)^2 \frac{10^4}{g}. \quad (2.17)$$

Besides the above conditions, the ratio r at the decay is bounded from above according to

$$r(T_d) < \frac{\sqrt{mM_p}}{T_d} \frac{m_\chi(gm\phi_{osc})^{3/2}}{m^2 M_p^2 (2\pi)^3} < \frac{\sqrt{mM_p}}{T_d} \frac{\phi_{osc}^2}{M_p^2}, \quad (2.18)$$

which is the usual condition that appears for the conventional curvatons. Since the curvaton decay rate Γ_χ will be at least of order m_χ^3/M_p^2 , there is a lower bound for T_d given by

$$T_d \sim \sqrt{\Gamma M_p} \geq \frac{m_\chi^{3/2}}{M_p^{1/2}}. \quad (2.19)$$

These conditions lead to

$$r(T_d) < 10^{-2} \times \frac{g^{3/2} \phi_{osc}^{3/2}}{M_p \sqrt{m_\chi^{bare}}} \simeq 10^4 \times \frac{m^{3/2}}{\sqrt{m_\chi^{bare}} M_p} \left| \frac{\sin 2\theta_0}{\cos 2\theta_0} \right|^3, \quad (2.20)$$

where we assumed that the curvaton decays when m_χ^{bare} dominates m_χ . Therefore, our conclusion in this section is that one can expect $r(T_d) \sim 1$ and the generation of the curvature perturbation from the hybrid-type curvaton.⁷

In the next section we would like to see what happens in the opposite limit of the broken symmetry, which is denoted by $x \gg 1$.

3 Badly broken symmetry ($x \gg 1$)

In the previous section we considered the hybrid-type curvaton that is generated through inhomogeneous preheating with the potential with $x \ll 1$. Alternatively, one may consider the opposite limit $x \gg 1$, where the scalar ϕ_2 becomes very light compared with ϕ_1 . Then, the oscillation will be induced solely by the “massive” field ϕ_1 , while the fluctuation is

⁷Alternatively, one may consider that the model helps to generate smaller-scale perturbations that could be highly non-Gaussian. This is an another possibility but it will not be discussed in this paper.

dominantly induced by the “light” field ϕ_2 . In this limit, the two roles (oscillation and fluctuation) of the field ϕ are distributed to the components ϕ_1 and ϕ_2 , which should be considered as independent fields. In this section the value of the dimensionless parameter is supposed to be about $x \sim 100$.⁸

In the limit $x \gg 1$, we cannot use the $\mathcal{O}(x)$ expansion that was used in Eq.(2.7). In this limit the spectrum of the perturbations is given by $H_I/2\pi$ for both $\delta\phi_1$ and $\delta\phi_2$, while $\dot{\delta\phi}_1$ and $\dot{\delta\phi}_2$ are given by $mH_I/2\pi$ and 0, respectively. Then at the minimum of the trajectory, we obtain $|\dot{\phi}_*| \simeq m\phi_{1,osc}$ and $|\phi_*| \simeq \phi_{2,osc}$. Therefore, the fluctuation of the number density is given by

$$\begin{aligned}\frac{\delta n_\chi}{n_\chi} &\simeq \left(\frac{3}{2} + \frac{\pi g |\phi_*|^2}{|\dot{\phi}_*|} \right) \frac{\dot{\delta\phi}_*}{|\dot{\phi}_*|} - \frac{2\pi g |\phi_*|^2}{|\dot{\phi}_*|} \frac{\delta\phi_*}{|\phi_*|} \\ &\simeq \left(\frac{3}{2} + \frac{\pi g |\phi_{2,osc}|^2}{|m\phi_{1,osc}|} \right) \frac{2\delta\phi_1}{|\phi_{1,osc}|} - \frac{2\pi g |\phi_{2,osc}|^2}{|m\phi_{1,osc}|} \frac{\delta\phi_2}{|\phi_{2,osc}|} \\ &\simeq \frac{3\delta\phi_1}{|\phi_{1,osc}|} - \frac{2\pi g |\phi_{2,osc}| \delta\phi_2}{m |\phi_{1,osc}|},\end{aligned}\quad (3.1)$$

where we assumed $\phi_{2,osc} \ll \phi_{1,osc}$. We will disregard the first term assuming that m is comparably smaller than $|\phi_{2,osc}|$. Then, the required condition for the COBE normalization is

$$\frac{g H_I}{m} \left| \frac{\phi_{2,osc}}{\phi_{1,osc}} \right| \sim 10^{-5}. \quad (3.2)$$

Note that in the scenario with slightly broken symmetry, the second term that is proportional to $\delta\phi_*$ was negligible; while in the present case the second term dominates the fluctuation. Because of the exponential factor in the form of n_χ , the required condition for the successful preheating is

$$\frac{\pi g |\phi_{2,osc}|^2}{m |\phi_{1,osc}|} < 1. \quad (3.3)$$

From Eq.(3.2) and (3.3), we find

$$\frac{\pi |\phi_{2,osc}|}{H_I} < 10^5. \quad (3.4)$$

The non-Gaussian parameter f_{NL} is of course different from the $x \ll 1$ case, since in the present case the term proportional to $\delta\phi_2$ will dominate the fluctuation. Since the expression for $\delta n_\chi/n_\chi$ up to the second order is given by

$$\frac{\delta n_\chi}{n_\chi} \simeq -\frac{2\pi g |\phi_{2,osc}| \delta\phi_2}{|\dot{\phi}_*|} - \frac{1}{2} \left(\frac{2\pi g}{|\dot{\phi}_*|} - \frac{4\pi^2 g^2 |\phi_{2,osc}|^2}{|\dot{\phi}_*|^2} \right) (\delta\phi_{2,osc})^2, \quad (3.5)$$

⁸Note that an extremely light scalar field will induce another cosmological problem to the theory.

the non-Gaussian parameter is

$$-\frac{3}{5}f_{NL} \simeq \frac{3|\dot{\phi}_*|}{4\pi g|\phi_{2,osc}|^2} - \frac{3}{2} \simeq \frac{3m|\phi_{1,osc}|}{4\pi g|\phi_{2,osc}|^2}. \quad (3.6)$$

Therefore, the condition $|f_{NL}| < 100$ is expressed as

$$\frac{m|\phi_{1,osc}|}{\pi g|\phi_{2,osc}|^2} < 10^2. \quad (3.7)$$

Combined with Eq.(3.2), we obtain

$$10^3 < \frac{\pi|\phi_{2,osc}|}{H_I}. \quad (3.8)$$

As a result, we find the initial condition for the light field ϕ_2 as

$$10^3 < \frac{\pi|\phi_{2,osc}|}{H_I} < 10^5. \quad (3.9)$$

As in the slightly broken case, the ratio r at the decay is bounded from above by

$$r(T_d) < 10^{-2} \times \frac{g^{3/2}\phi_{1,osc}^{3/2}}{M_p\sqrt{m_\chi^{bare}}} \quad (3.10)$$

which leads to the initial condition $(g\phi_{1,osc})^3 > 10^4 M_p^2 m_\chi^{bare}$. Therefore, generation of the curvature perturbation from the hybrid-type curvaton is possible for $x \gg 1$.

4 Conclusions and Discussions

In this paper we have considered inhomogeneous preheating that is characterized by a potential with broken symmetry, and looked to see if this could be used to generate curvatons. As we have shown in this paper, this simple realization of the hybrid of the two seemingly different ideas is successful.

To complement the discussion, we will summarize the situation in this paper and compare it with other extended models. First, there is the possibility that the preheat field decays instantaneously after preheating. This is the so-called instant preheating scenario[19] which has been used to generate the cosmological perturbations in Ref. [13] for $x \ll 1$, and in Ref. [14, 15] for $x \gg 1$. Instead of considering the instant decay, we considered in this paper a long-lived curvaton generation. Combined with the curvatons, the inhomogeneous preheating scenario opens another possibility for generating the cosmological perturbations from late-time oscillation.

We used two scalar fields ϕ_1 and ϕ_2 . Our model is characterized by the ϕ_2 -potential, which is shown in Fig.2 and Fig.3.

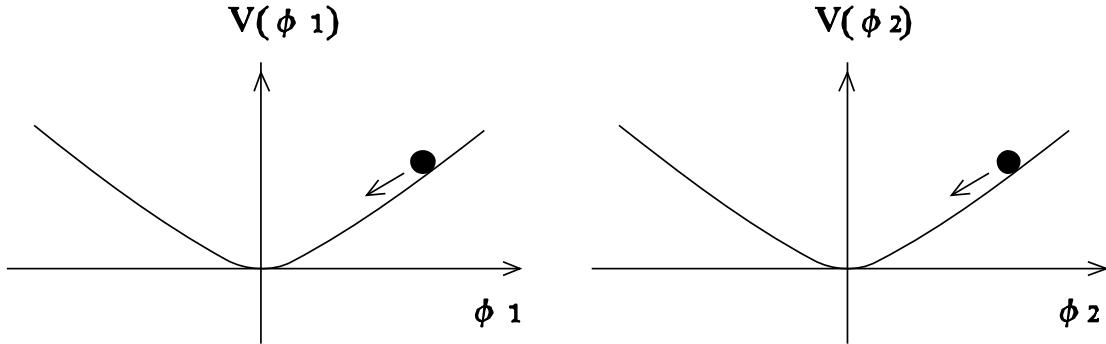


Figure 2: This figure shows the potential for slightly broken symmetry ($x \ll 1$). ϕ_1 and ϕ_2 start to oscillate at the same time.

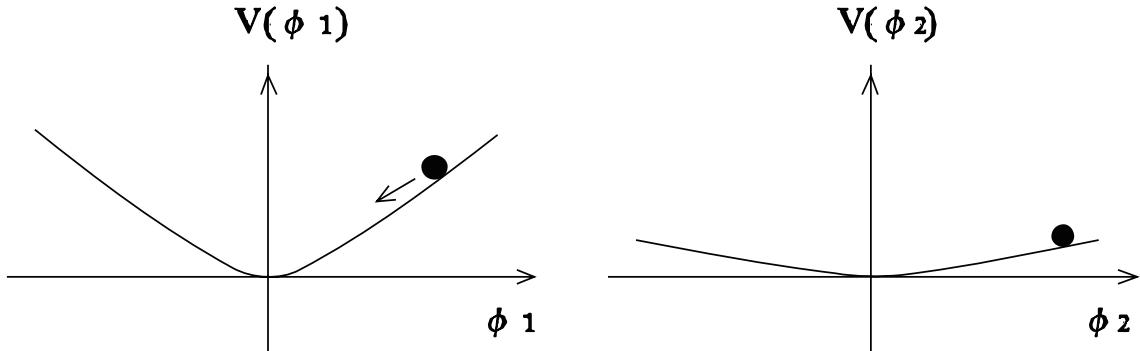


Figure 3: This figure shows the potential for badly broken symmetry ($x \gg 1$). ϕ_1 oscillates to induce preheating, while ϕ_2 induces m_χ -perturbation to the preheating.

Let us explain what happens using the potential depicted in Fig.4. Besides the hybrid curvatons we considered in this paper, one may consider an alternative to the curvatons. Namely, trapped inflation is realized by the ϕ_2 -potential that is depicted in Fig.4, which has a potential similar to thermal inflation. Instead of considering thermal trapping, we considered in Ref. [20] the trapping induced by the preheat field. This scenario is based on the trapped inflation[21] that has originally been discussed with string model, but the situation discussed in Ref. [20] is rather different from the original string model.

Finally, we will comment on the possibility that ϕ_2 may have quintessential potential, which is shown in Fig.5. This model has been discussed previously by us in Ref. [17]. We showed that the curvaton scenario is successful in this model. Moreover, this model can

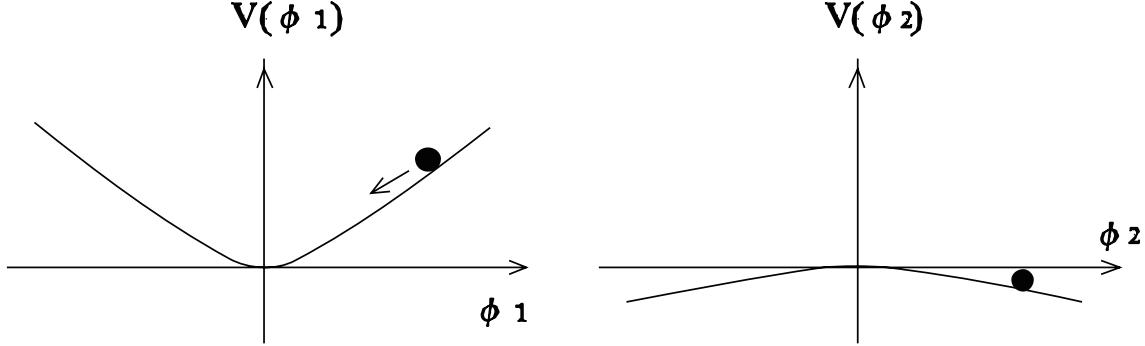


Figure 4: This figure shows the potential for a trapped inflationary model that has been discussed in Ref. [20]. In this case the primordial isocurvature perturbation $\delta\phi_2$ is converted into δn_χ through the inhomogeneous preheating, which leads to the fluctuation of the number of e-foldings δN_e . See Ref. [9] for more details of the mechanism for generating the curvature perturbation at the end of inflation.

be regarded as the hybrid of chaotic-type inflation and quintessential inflation connected at the ESP.

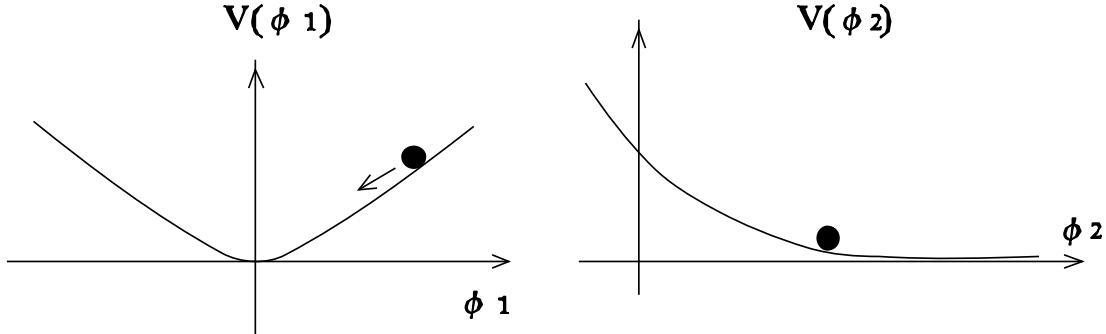


Figure 5: This picture shows the potential for hybrid-type quintessential inflaton that has been discussed in Ref. [17].

Now we would like to summarize the motivation and the results obtained in these successive approaches. In these seemingly different models we considered hybrids of inhomogeneous preheating and other cosmological scenarios (such as curvatons, trapped and quintessential inflation) to obtain new alternatives to the traditional inflationary scenario. We know that preheating is a common phenomenon that is expected to occur at the end of inflation, or when a field starts to oscillate at a later epoch. It is also very natural to expect that the potential is characterized by broken symmetry. Motivated by these

observations, we considered in this paper the inhomogeneous preheating characterized by broken symmetry. For the instant-decay scenario, generation of the perturbations has been discussed in Ref. [13, 14, 15]. Our present model is different from these scenarios, as we considered in this paper the late-decay curvaton.

5 Acknowledgment

We wish to thank K.Shima for encouragement, and our colleagues at Tokyo University for their kind hospitality.

A Instant-decay scenario with slightly broken symmetry

In this appendix we will explain how the fluctuation related to the number density of the preheat field is generated from the oscillation using the potential characterized by slightly broken symmetry. Although what we will explain here has already been discussed in Ref. [13], we think it is important to review the issue from the viewpoint of applying the inhomogeneous preheating to the curvaton generation. It is also important to examine the required condition in the original model and explain how it could be removed or relaxed in the hybrid model.

Here we will consider the potential of the scalar field described by 2 degrees of freedom (ϕ_1 and ϕ_2)

$$V(\phi_1, \phi_2) = \frac{m^2}{2} \left[\phi_1^2 + \frac{\phi_2^2}{1+x} \right], \quad (\text{A.1})$$

where x represents a measure of the symmetry breaking. Since we are considering slightly broken symmetry $x \ll 1$, it would be useful to introduce a complex field ϕ that is defined as

$$\phi = \phi_1 + \phi_2 = |\phi| e^{i\theta}. \quad (\text{A.2})$$

The oscillation of the scalar field ϕ will induce preheating for the preheat field χ if there is interaction given by $\mathcal{L} = -\frac{1}{2}g^2|\phi|^2\chi^2$. In the original scenario the preheat field was

assumed to have a negligible mass term, $m_\chi^{bare} \simeq 0$.⁹ Although the bare mass is negligible, there is the effective mass $m_\chi(\phi)$ that depends on the time-dependent value of ϕ . Since the preheat field becomes massless near the enhanced symmetric point (ESP) at $|\phi| = 0$, χ will be generated near the ESP. Then the comoving number density of the preheat field χ produced at the first scattering is given by

$$n_\chi^{(1)} = \frac{(g|\dot{\phi}_*|)^{3/2}}{(2\pi)^3} \exp \left[-\frac{\pi g |\phi_*|^2}{|\dot{\phi}_*|} \right], \quad (\text{A.3})$$

where $*$ denotes the value when the oscillating field reaches the minimum of the potential along its trajectory. Expressing the initial conditions as $|\phi_0|$ and θ_0 , these values are given by

$$|\phi_*| \simeq \frac{|\phi_0| \pi x}{2\sqrt{2}} |\sin 2\theta_0| \quad (\text{A.4})$$

$$|\dot{\phi}_*| \simeq m|\phi_0| \sqrt{1 - x \sin^2 \theta_0}. \quad (\text{A.5})$$

Considering the primordial perturbation $\delta\theta_0$, one obtains

$$\frac{\delta|\phi_*|}{|\phi_*|} \simeq \frac{2 \cos 2\theta_0}{\sin 2\theta_0} \delta\theta_0 - (\delta\theta_0)^2 \quad (\text{A.6})$$

$$\frac{\delta|\dot{\phi}_*|}{|\dot{\phi}_*|} \simeq -\frac{x}{2} [\sin 2\theta_0 \delta\theta_0 + \cos 2\theta_0 (\delta\theta_0)^2]. \quad (\text{A.7})$$

Then the fluctuation that could be induced for the number density of the preheat field is given by

$$\frac{\delta n_\chi}{n_\chi} \simeq -\frac{3x}{4} [\sin 2\theta_0 \delta\theta_0 + \cos 2\theta_0 (\delta\theta_0)^2]. \quad (\text{A.8})$$

where we did not consider terms proportional to x^2 , since they are very small in the limit $x \ll 1$. Because of the exponential factor appearing in the form of n_χ , the condition that is needed for successful preheating is

$$\frac{\pi g |\phi_*|^2}{|\dot{\phi}_*|} \simeq \frac{\pi g (\phi_0 x \sin 2\theta_0)^2}{m \phi_0} < 1, \quad (\text{A.9})$$

which suggests that $x^2 < m/(g|\phi_0| \sin^2 2\theta_0)$. Since in the original scenario ϕ is assumed to be the inflaton of a chaotic-type inflationary model, there is an upper bound for m

⁹Please note that this is not true in our hybrid curvaton model.

$(m < 10^{13} GeV)$, so that the traditional perturbation generated by the inflaton does not dominate the primordial perturbation. As a result, x must satisfy the condition

$$x^2 < 10^{-5} g^{-1} \sin^{-2} 2\theta \quad (\text{A.10})$$

for $\phi_0 \simeq M_p$. Moreover, if the cosmological perturbation is generated by the inhomogeneous preheating, there is the condition

$$\left[\frac{3x \sin 2\theta_0}{4} \right] \delta\theta_0 \sim 10^{-5}, \quad (\text{A.11})$$

which leads to

$$x \sim 10^{-5} \frac{4}{(3 \sin 2\theta_0) \delta\theta_0}. \quad (\text{A.12})$$

Besides the above condition, there is another condition coming from the non-Gaussian parameter $f_{NL} < 100$, which is

$$x > 0.01 \frac{\cos 2\theta_0}{\sin^2 2\theta_0}. \quad (\text{A.13})$$

In order to satisfy Eq.(A.10) and (A.13), the coupling constant must satisfy the condition

$$g < 10^{-5} \left(\frac{\hat{f}_{NL} \sin 2\theta_0}{\cos 2\theta_0} \right)^2, \quad (\text{A.14})$$

where \hat{f}_{NL} denotes the upper bound for f_{NL} . According to the original analysis[19], the ratio of the energy density of the preheat field to the energy density of the inflaton field will be $r_0 \equiv \frac{\rho_\chi}{\rho_\phi} \sim 2g^{5/2}$, which suggests

$$r_0 < 10^{-25/2} \left(\frac{\hat{f}_{NL} \sin 2\theta_0}{\cos 2\theta_0} \right)^5, \quad (\text{A.15})$$

while $r_0 \sim 1$ is required if the cosmological perturbation is supposed to be generated by the preheat field.

Besides the above example which is characterized by $x \ll 1$, we considered in Ref. [14, 15] a new model that is characterized by $x \gg 1$, where the lighter field ϕ_2 is almost independent of the oscillating field.

Unlike the model discussed in Ref. [13, 14, 15] in which the preheat field is assumed to decay instantly after preheating, we will “not” assume in this paper that the preheating is accompanied by instant decay. We will assume that the decay rate of the preheat

field is rather small so that the preheat field decays late. A small decay rate is due to small couplings to light fields. This gives a complementary scenario for the instant decay scenario. Assuming that ϕ starts to roll when $m \sim H$, the required condition for the stability of the preheat field at the maximum value of the oscillating field is

$$g_d^2 \times g < \frac{m}{\phi_{max}}, \quad (\text{A.16})$$

where we assumed that the decay rate is given by $\Gamma_\chi = g_d^2 \times g \phi_{max}$, where ϕ_{max} ($\phi_{max} \ll \phi_{osc}$) denotes the maximum value of the oscillating field after the first scattering.¹⁰ Since the preheat field does not decay instantaneously after preheating, it evolves like a massive matter field when it has a bare mass term. Since the ratio of the massive matter to the radiation energy density increases during the radiation-dominated era, we do not have to assume that the density of the preheat field dominates the Universe just after preheating. This removes the condition $r_0 \sim 1$. As a result, the oscillating field that induces preheating is not necessarily the inflaton. Remember that in the Affleck-dine mechanism the late-time oscillation induced by a light scalar field plays an important role in generating the baryon asymmetry of the Universe. We showed that the generation of the perturbed curvaton is possible even if the inhomogeneous preheating is induced by late-time oscillation.

References

- [1] D. H. Lyth, *Constraints on TeV scale hybrid inflation and comments on non-hybrid alternatives*, *Phys.Lett.B* **466** (1999) 85 [hep-ph/9908219]; P. Kanti and K. A. Olive, *Assisted chaotic inflation in higher dimensional theories*, *Phys. Lett. B* **464** (1999) 192 [hep-ph/9906331]; N. Arkani-Hamed, S. Dimopoulos, N. Kaloper, and J. March-Russell, *Rapid asymmetric inflation and early cosmology in theories with submillimeter dimensions*, *Nucl.Phys.B* **567** (2000) 189 [hep-ph/9903224]; T. Matsuda, *Kaluza-Klein modes in hybrid inflation*, *Phys.Rev.D* **66** (2002) 107301 [hep-ph/0209214]; R. N. Mohapatra, A. Perez-Lorenzana, and C. A. de S.

¹⁰ ϕ_{max} is much smaller than ϕ_{osc} because of the confining potential induced by $n_\chi \neq 0$. See Ref. [21] for more details.

Pires, *Inflation in models with large extra dimensions driven by a bulk scalar field*, *Phys.Rev.D*62(2000)105030 [hep-ph/0003089]; A. Mazumdar, *Extra dimensions and inflation*, *Phys.Lett.B*469(1999)55 [hep-ph/9902381]; A. M. Green and A. Mazumdar, *Dynamics of a large extra dimension inspired hybrid inflation model*, *Phys.Rev.D*65(2002)105022 [hep-ph/0201209]; A. Mazumdar, *New developments on embedding inflation in gauge theory and particle physics* [arXiv:0707.3350]; T. Matsuda, *Nontachyonic brane inflation*, *Phys.Rev.D*67(2003)083519 [hep-ph/0302035]; T. Matsuda, *F term, D term and hybrid brane inflation*, *JCAP* 0311(2003)003 [hep-ph/0302078]; T. Matsuda, *Successful D term inflation with moduli*, *Phys.Lett.B*423(1998)35 [hep-ph/9705448].

- [2] A. R. Liddle and D. H. Lyth, *Cosmological inflation and large-scale structure*, Cambridge, UK: Univ. Pr. (2000).
- [3] A. Masiero, M. Peloso, L. Sorbo, and R. Tabbash, *Baryogenesis versus proton stability in theories with extra dimensions*, *Phys.Rev.D*62(2000)063515 [hep-ph/0003312]; T. Matsuda, *Baryon number violation, baryogenesis and defects with extra dimensions*, *Phys.Rev.D*66(2002)023508 [hep-ph/0204307]; T. Matsuda, *Enhanced baryon number violation due to cosmological defects with localized fermions along extra dimensions*, *Phys.Rev.D*65(2002)107302 [hep-ph/0202258].
- [4] R. Allahverdi, K. Enqvist, A. Mazumdar and A. Perez-Lorenzana, *Baryogenesis in theories with large extra spatial dimensions*, *Nucl.Phys. B*618(2001)277 [hep-ph/0108225]; T. Matsuda, *Affleck-Dine baryogenesis in the local domain*, *Phys.Rev.D*65(2002)103502 [hep-ph/0202211]; T. Matsuda, *Affleck-Dine baryogenesis after thermal brane inflation*, *Phys.Rev.D*65(2002)103501 [hep-ph/0202209].
- [5] G. R. Dvali, G. Gabadadze, *Nonconservation of global charges in the brane universe and baryogenesis*, *e Phys.Lett.B*460(1999)47 [hep-ph/9904221]; A. Pilaftsis, *Leptogenesis in theories with large extra dimensions*, *Phys.Rev.D*60(1999)105023 [hep-ph/9906265]; S. Davidson, M. Losada, and A. Riotto, *A new perspective on baryogenesis*, *Phys.Rev.Lett.*84(2000)4284 [hep-ph/0001301]; T. Mat-

suda, *Electroweak baryogenesis mediated by locally supersymmetry breaking defects*, *Phys.Rev.D*64(2001)083512 [hep-ph/0107314].

- [6] A. D. Linde and V. F. Mukhanov, *Nongaussian isocurvature perturbations from inflation*, *Phys.Rev.D*56(1997)535 [astro-ph/9610219]; D. H. Lyth, D. Wands, *Generating the curvature perturbation without an inflaton*, *Phys.Lett.B*524(2002)5 [hep-ph/0110002]; T. Moroi, T. Takahashi, *Effects of Cosmological Moduli Fields on Cosmic Microwave Background*, *Phys.Lett.B*522(2001)215 [hep-ph/0110096].
- [7] D. H. Lyth, *Can the curvaton paradigm accommodate a low inflation scale?*, *Phys.Lett.B*579(2004)239-244,2004[hep-th/0308110].
- [8] T. Matsuda, *Curvaton paradigm can accommodate multiple low inflation scales*, *Class.Quant.Grav.*21(2004)L11 [hep-ph/0312058]; K. Dimopoulos, D. H. Lyth and Y. Rodriguez, *Low scale inflation and the curvaton mechanism*, *JHEP* 0502:055,2005 [hep-ph/0411119]; K. Dimopoulos and G. Lazarides, *Modular inflation and the orthogonal axion as curvaton*, *Phys.Rev.D*73(2006)023525[arXiv:hep-ph/0511310]; K. Dimopoulos, *Inflation at the TeV scale with a PNGB curvaton*, *Phys.Lett.B*634(2006)331 [arXiv:hep-th/0511268]; K. Dimopoulos.
- [9] D. H. Lyth, *Generating the curvature perturbation at the end of inflation*, *JCAP* 0511:006,2005 [astro-ph/0510443]; F. Bernardeau, L. Kofman and J.P. Uzan, *Modulated fluctuations from hybrid inflation*, *Phys.Rev.D*70(2004)083004 [astro-ph/0403315]; T. Matsuda, *Brane inflation without slow-roll*, *JHEP* 0703:096,2007 [astro-ph/0610402].
- [10] F. Vernizzi, *Generating cosmological perturbations with mass variations*, *Nucl.Phys.Proc.Suppl.*148(2005)120 [astro-ph/0503175]; F. Vernizzi, *Cosmological perturbations from varying masses and couplings*, *Phys.Rev.D*69(2004)083526 [astro-ph/0311167].
- [11] G. Dvali, A. Gruzinov and M. Zaldarriaga, *A new mechanism for generating density perturbations from inflation*, *Phys.Rev.D*69(2004)023505 [astro-ph/0303591]; G. Dvali, A. Gruzinov and M. Zaldarriaga, *Cosmological perturbations from inho-*

mogeneous reheating, freezeout, and mass domination, *Phys.Rev.D*69(2004)083505 [astro-ph/0305548].

- [12] K. Enqvist, A. Mazumdar and M. Postma, *Challenges in generating density perturbations from a fluctuating inflaton coupling*, *Phys.Rev.D*67(2003)121303 [astro-ph/0304187]; A. Mazumdar and M. Postma, *Evolution of primordial perturbations and a fluctuating decay rate*, *Phys.Lett.B*573(2003), [*Erratum-ibid. B*585(2004)295] [arXiv:astro-ph/0306509]; S. Tsujikawa, *Cosmological density perturbations from perturbed couplings*, *Phys.Rev.D*68(2003)083510 [astro-ph/0305569]; B. A. Bassett, S. Tsujikawa and D. Wands, *Inflation dynamics and reheating*, *Rev.Mod.Phys.*78(2006)537 [astro-ph/0507632]; R. Allahverdi, K. Enqvist, A. Jokinen and A. Mazumdar, *Identifying the curvaton within MSSM*, *JCAP*0610:007,2006 [hep-ph/0603255]; R. Allahverdi, *Scenarios of modulated perturbations*, *Phys.Rev.D*70(2004)043507 [astro-ph/0403351];
- [13] E. W. Kolb, A. Riotto and A. Vallinotto, *Curvature Perturbations from Broken Symmetries*, *Phys.Rev.D*71(2005)043513[astro-ph/0410546].
- [14] T. Matsuda, *Generating the curvature perturbation with instant preheating*, *JCAP*0703:003,2007 [hep-th/0610232].
- [15] T. Matsuda, *Generating curvature perturbations with MSSM flat directions*, *JCAP*0706:029,2007[hep-ph/0701024].
- [16] T. Matsuda, *Topological curvatons*, *Phys.Rev.D*72(2005)123508 [hep-ph/0509063].
- [17] T. Matsuda, *NO Curvatons or Hybrid Quintessential Inflation*, *JCAP*08(2007)003[arXiv:0707.1948].
- [18] J.C. Bueno Sanchez and K. Dimopoulos, *Trapped Quintessential Inflation*, *Phys.Lett.B*642(2006)294, *Erratum-ibid.B*647(2007)526 [hep-th/0605258]; J.C. Bueno Sanchez and K. Dimopoulos, *Trapped quintessential inflation in the context of flux compactifications* [hep-th/0606223].
- [19] G. N. Felder, L. Kofman, A. D. Linde, *Instant preheating*, *Phys.Rev.D*59(1999)123523 [hep-ph/9812289].

- [20] T. Matsuda, *Cosmological perturbations from inhomogeneous preheating and multi-field*, *JHEP* 0707:035, 2007 [arXiv:0707.0543].
- [21] L.Kofman, A. Linde, X. Liu, A. Maloney, L. McAllister and E. Silverstein *Beauty is attractive: Moduli trapping at enhanced symmetry points*, *JHEP* 0405:030, 2004 [hep-th/0403001]; S. Watson, *Moduli stabilization with the string Higgs effect*, *Phys. Rev. D* 70 (2004) 066005 [hep-th/0404177].